The H.E.S.S. experiment - Recent Results and Future

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LLR Ecole Polytechnique

IWFM, Habère-Poche, February 2005

- The H.E.S.S.detector (Phase 1)
- Galactic Sources
- Extragalactic Sources
- The future: Phase 2
- Summary

Sources seen by H.E.S.S. ($S > 5\,\sigma$)

Object	Flux level [Crab]	First detection	Conf.	Contradiction
Crab Nebula	1.00	Whipple	Many	–
RXJ 1713.7-3946	0.70	CANGAROO	HESS	Spectrum
Vela Junior	?	CANGAROO	HESS	??
G0.9-0.1	0.02	HESS		
Sgr A*	0.1-0.4	CANGAROO	HESS	Spectrum
HESS J1303-63	0.10	HESS	–	–
PSR B1259-63/SS2823	0.05	HESS	–	–
Mkn 421	0.2->1	Whipple	Many	–
PKS 2155-304	0.1-0.6	Durham	HESS	–

... more to come soon ...

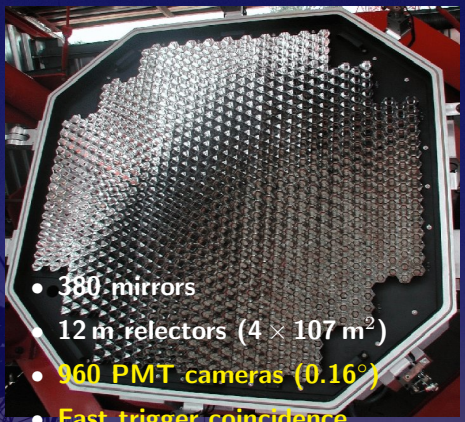
no detection of :

SN 1006, Vela, PSRB1706-44

The H.E.S.S.Detector

HESS 1: Stereoscopic System of 4 Cherenkov Telescopes

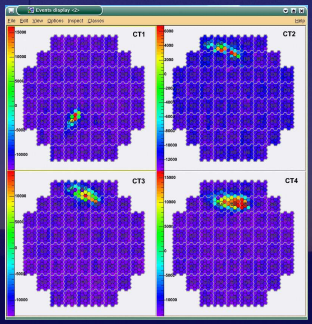
- Namibia, Khomas Highlands (23°S, 15°E)
- Altitude 1800 m



- 380 mirrors
- 12 m relectors ($4 \times 107\,m^2$)
- 960 PMT cameras (0.16°)
- Fast trigger coincidence
- Electronics integrated in Cameras
- 5.3° FoV
- Energy threshold 100 GeV

HESS I – Phase 1 Completed

First light telescope 1	:	June 2002
Two telescopes	:	March 2003
Stereoscopy	:	July 2003
Three telescopes	:	September 2003
Four telescopes	:	December 2003

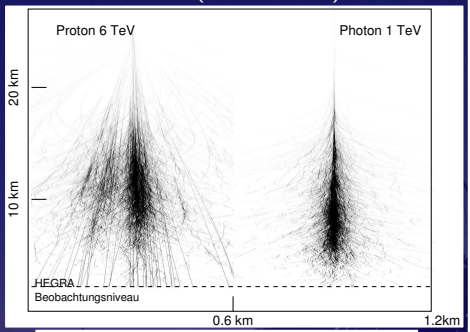


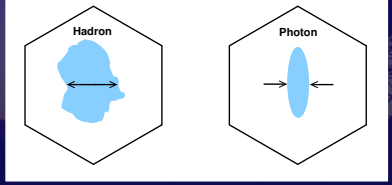
Phase 1 completed & fully operational !

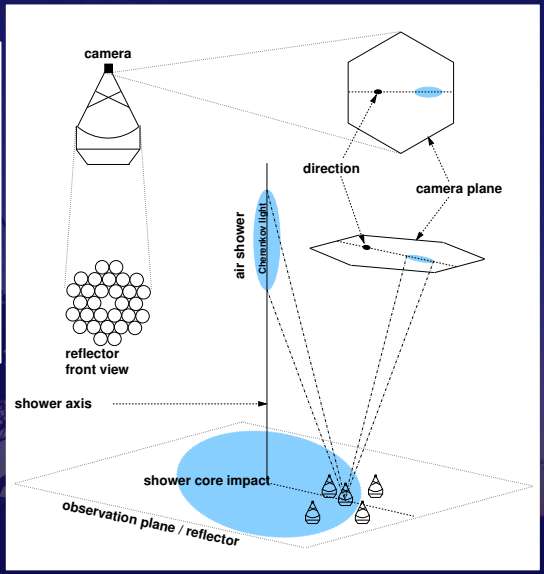
- 0.01 Crab in 25 h
- E_{thr} = 100 GeV (max. diff. E-spec.)

Imaging Air Shower Cherenkov Technique (Pioneered by Whipple)

MC simulation (Horns 2000):

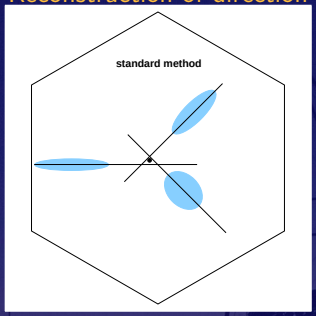






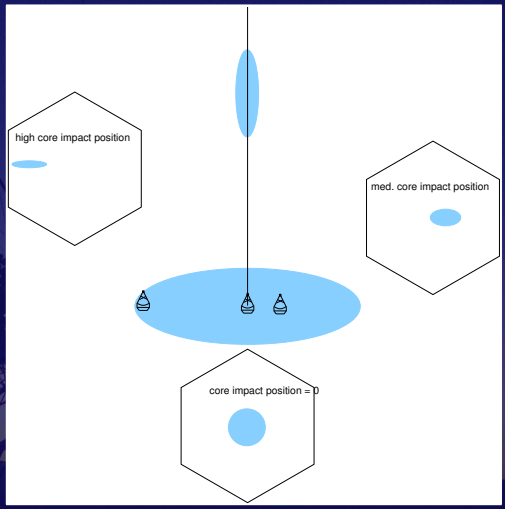
The Basic Stereoscopic Reconstruction Principles

Reconstruction of direction

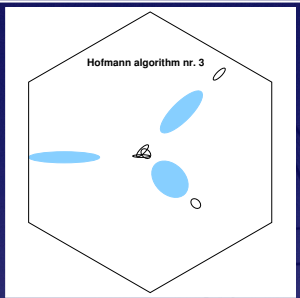


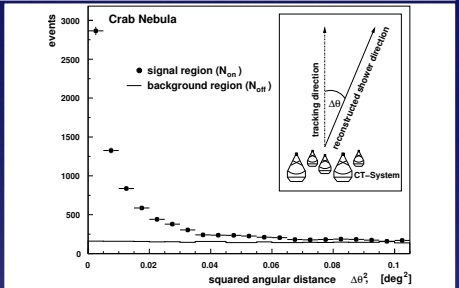
Scaling the image width: Mean Scaled Width Parameter

- Reconstruct: $core$ + amplitude + zenith angle
- MC lookup tables for width $\longrightarrow$ scaling of the width



Stereoscopic Observation Technique - Pioneered by HEGRA

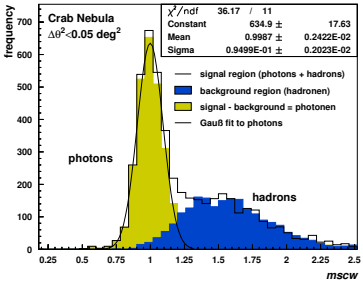




$\geq$ 2 images $\rightarrow$ superposition:

- ... in the camera $\rightarrow$ direction (θ)
- ... at observation level
  $\rightarrow$ core impact position

Hadron rejection:
$core$ + amplitude + zenith angle
$\rightarrow$ $m$ean scaled $w$idth ($mscw$)



Simulation, Calibration & Analysis

Monte Carlo simulations

- CORSIKA + sim_hessarray
- KASKADE + smash

2 Calibration chains

- Heidelberg
- Paris

Background subtraction

- Geometric models
- Template-model
- Likelihood-based model

Different Calibration methods

- Single ph.e.
- Muon rings
- Laser System

Shower reconstruction methods

- Standard Hillas reconstruction
- Semi-analytical model
- 3D-Model

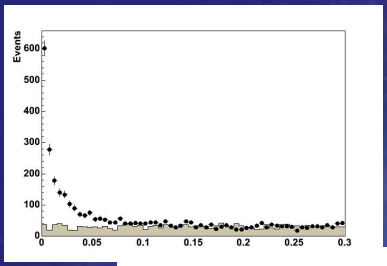
Redundance gives confidence: Robust results

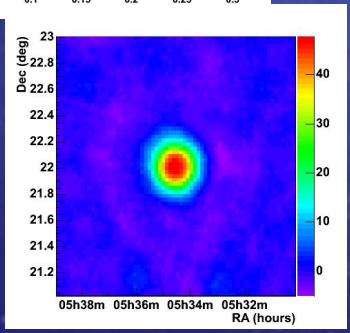
Galactic Sources

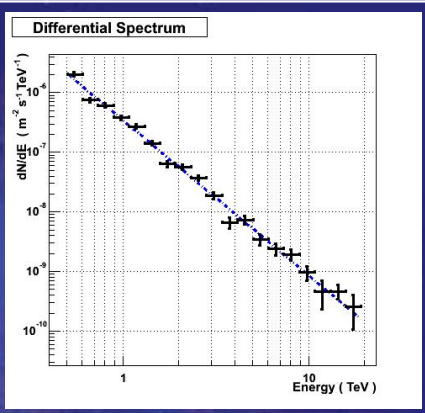
Physics: Origin/acceleration of Cosmic Rays, new sources, Dark matter ...

SNRs

The Crab Nebula as seen by HESS (preliminary)





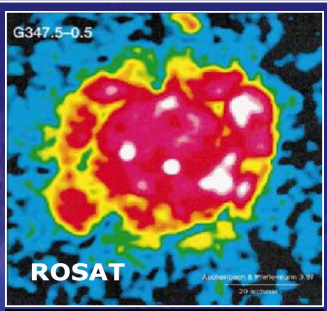


- High zenith angle (E_{thr}≈325 GeV)
- Independent analyses give consistent results
- Spectral index: $\alpha = 2.62 \pm 0.02$
- Compatible with previous results

The Shell-Type Supernova Remnant RXJ 1713-3946

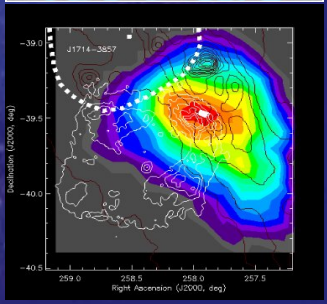
Discovery in X-rays

- ROSAT All-Sky survey source
- Non-thermal X-rays
- Distance: 1 kpc (CO survey)
- Angular extension: 1 deg



First TeV-detection: CANGAROO II
(Muraishi, A. et al. 2000; Enomoto, R. et al. 2002)

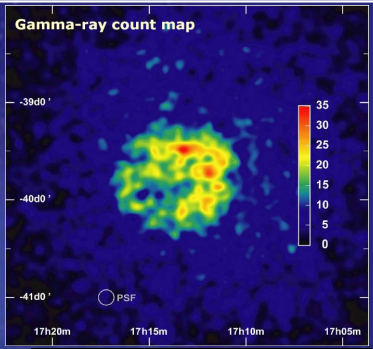
- ≈0.7 Crab
- Question of Cosmic ray acceleration ⟶ controversial discussions (Pohl et al. 2002)



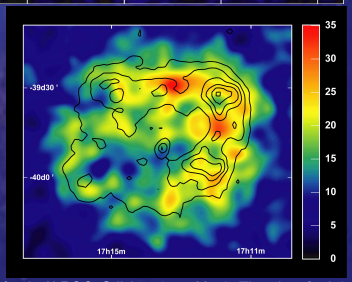
RXJ 1713.7-3946 as seen by HESS

Recent GeV/TeV-confirmation: HESS 2004
(D. Berge, Gamma 2004 & Nature acc. f. publ.)

- High quality data 18.1 h
- $>20\,\sigma$ total remnant
- *The first ever astronomical TeV-image*
- Shown here: High resolution data sub-sample (E>800GeV)



Superposition:
ASCA X-ray data contours



RXJ 1713.7-3946 – The spectrum

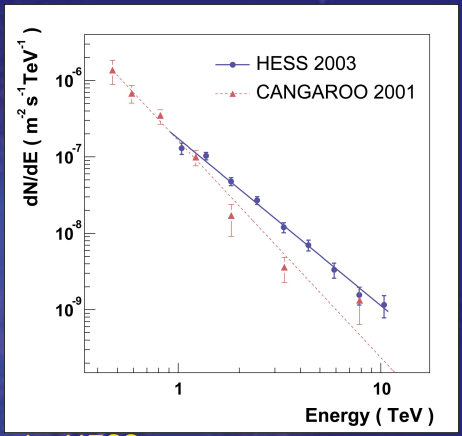
Spectrum: inconsistent spectral indices

H.E.S.S.:

- $\alpha = 2.19 \pm 0.09 \pm 0.15$
- consistent results in independent analyses

CANGAROO II:

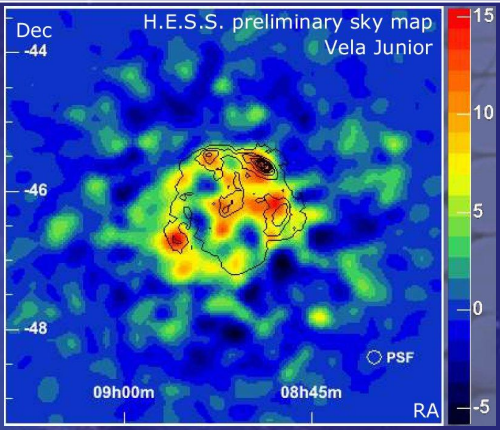
- $\alpha = 2.84 \pm 0.15 \pm 0.20$



$\longrightarrow$ soon: results of 2004 observations by HESS
$\longrightarrow$ GLAST observations of 70 MeV bump ?

Vela Junior

- shell-type SNR, diameter 2 deg
- superimposed with Vela
- distance 200 pc - 2 kpc ?
- non-thermal component of X-ray emission (ROSAT, ASCA, XMM)
- detection by CANGAROO



H.E.S.S.-observations

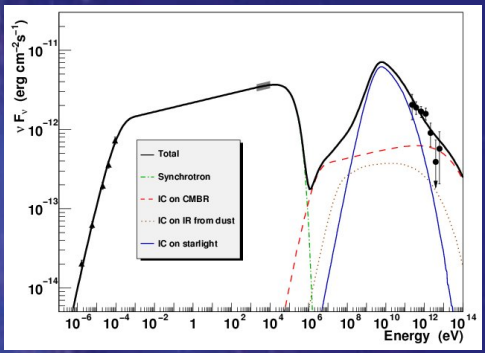
- 3.2 h 4-telescope data at 25 deg
- hard cuts : 675 GeV threshold
- bg : 6.8 h off source data

Texas Symposium, N. Komin, D. Berge, M. Lemoine-Goumard, M. de Naurois

- overall spectrum : $\gamma = 2.3 \pm 0.2$
- $\Phi(> 1TeV) = (2.0 \pm 0.5)10^{-11}$

G0.9-0.1

- *composite* SNR
- $\approx$ 0.13 deg diameter shell
- $\approx$ 0.03 deg core region: pulsar wind nebula



A&A submitted
origin of gamma-rays in core
well explained by IC

H.E.S.S.-observations

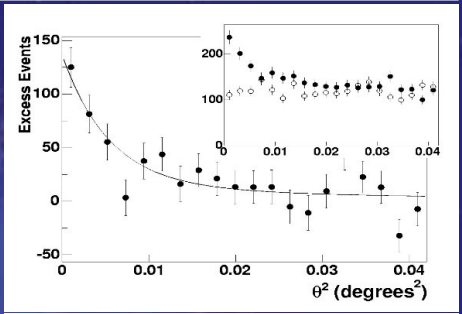
- 50 h at 18 deg Z.A.
- E_{thr} = 170 GeV, 13 σ
- $\Phi(> 200\,GeV) = (5.7 \pm 0.7 \pm 1.2)10^{-12}\mathrm{cm}^{-2}\mathrm{s}^{-1}$
- power law index γ = 2.4$\pm$0.11$\pm$0.20

A unique GeV/TeV emitter

The Galactic Center (Sgr A*): HESS data

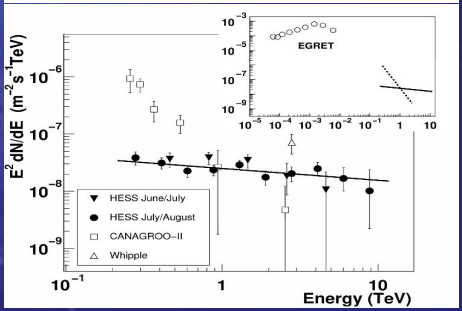
The Signal:

- 2-Telescope data
- Average zenith angle: 20°
- Two detector configurations
  - 4.7 h at $E_{thr} = 255$ GeV
  - 11 h at $E_{thr} = 165$ GeV
- Total significance: $9.2\,\sigma$

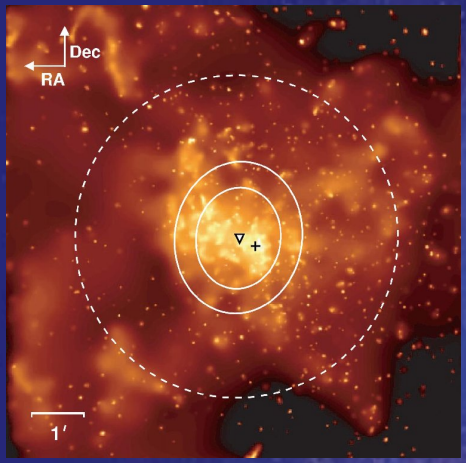


Spectrum:

- Power-law, $\alpha = 2.21 \pm 0.09 \pm 0.15$
- Steady state 0.05 % (E> 165 GeV)
- Strong contradiction to CANGAROO II



The Galactic Center (Sgr A*)



Chandra X-ray image

HESS superimposed:

- 68 % & 95 % confidence regions for source position
- 95 % upper limit on rms source size
- Position compatible within errors (30") with SgrA*
- $\Omega_{\mathbf{err}}$ reduced by 100 (as cmp to previous measurements)

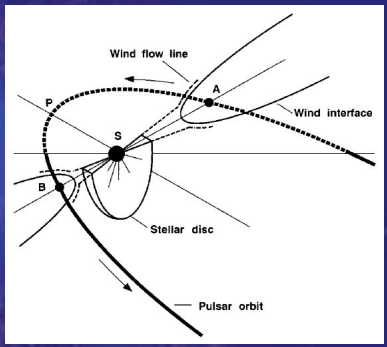
Dark Matter hypothesis: HESS spectrum + angular distribution $\Longrightarrow$ $M\chi > 12\,\mathrm{TeV}$ (90 % C.L.) (Horns, astro-ph/0408192, also: Aharonian & Neronov, astro-ph/0408303)

A new unique GeV/TeV emitter

PSRB 1259-63 / SS 2883

The Binary Pulsar system PSRB 1259-63 / SS 2883 at 1.5 kpc

- 10 $M_\odot$ Be star $L = 3\times10^{30}$ W
  dense stellar disk, high mass outflow
- 48 ms radio Pulsar
  $L_{spindown} = 8\times10^{28}$ W
- Pulsar orbit around Be star
  - 3.4 years
  - Periastron : 23 $R_\odot$
  - Apastron : 331 $R_\odot$
  - Inclination : 35 deg
  - Diameter: 350 $R_\odot$ (Point-like)

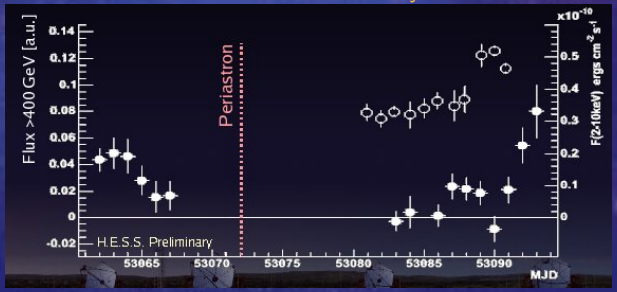


CANGAROO 3.8 m: 4.8 σ (1994), 10 m: Upper limits (2000), after periastron
HESS-observations at last periastron passage : 7^{th} of March 2004

PSRB 1259-63 / SS 2883 – HESS Results (S. Schlenker, Gamma 2004)

Pre-Periastron	Post-Periastron
• High quality data : 7.8 h	• High quality data : 17.4 h
• Significance : $9.1\,\sigma$	• Significance : $6.3\,\sigma$
• Excess rate : $\approx 0.4\,\gamma$/min	• Excess rate : $\approx 0.2\,\gamma$/min

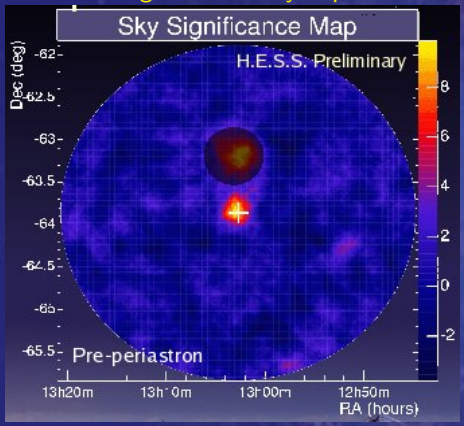
Overall $> 10\sigma$ detection by HESS



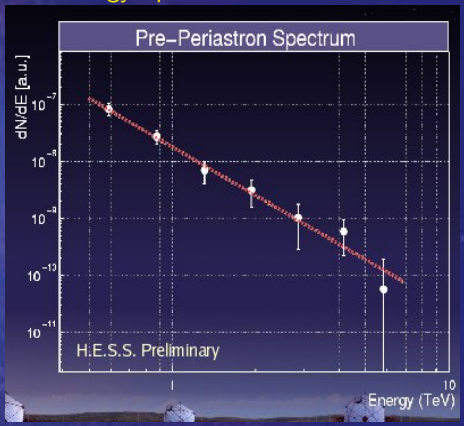
Flux E>400 GeV $\approx 5\,\%$ Crab

PSRB 1259-63 / SS 2883 – HESS Results (S. Schlenker Gamma 2004)

Significance Skymap



Energy Spectrum $\alpha = 2.8 \pm 0.3$



Unidentified GeV/TeV emitters

HESS J1303-63 (M. Beilicke, Gamma 2004)

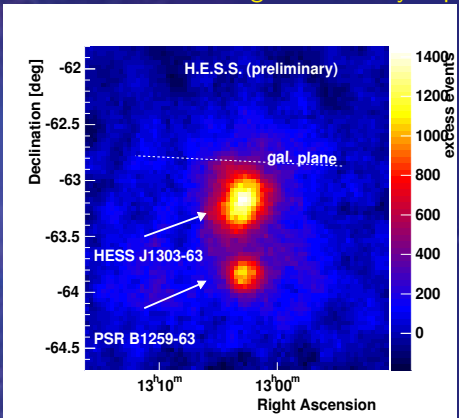
Surprise in PSRB 1259-63 observations:

- Second signal in FoV !
- 0.7 deg from Pulsar position
- Steady state signal ($18\,\sigma$)

Calibration or Physics ?

- High data quality
- Consistent results for different algorithms
- ...

Significance Skymap

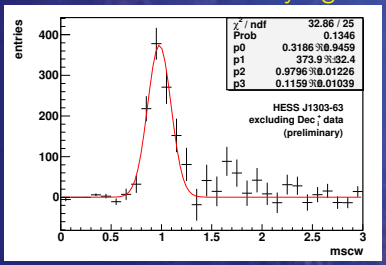


HESS J1303-63 (M. Beilicke, Gamma 2004)

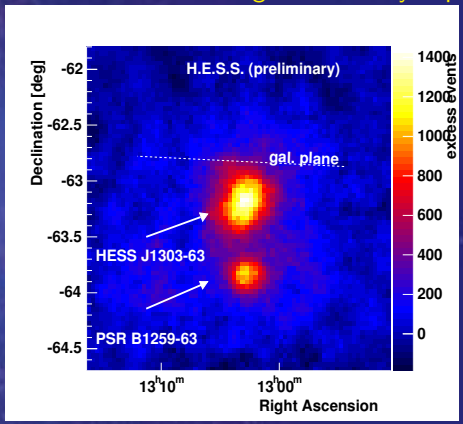
Surprise in PSRB 1259-63 observations:

- Second signal in FoV !
- 0.7 deg from Pulsar position
- Steady state signal (18σ)

HESS 1303-63: Gamma-ray signature

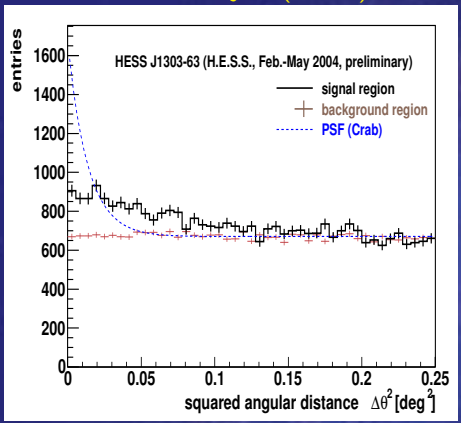


Significance Skymap

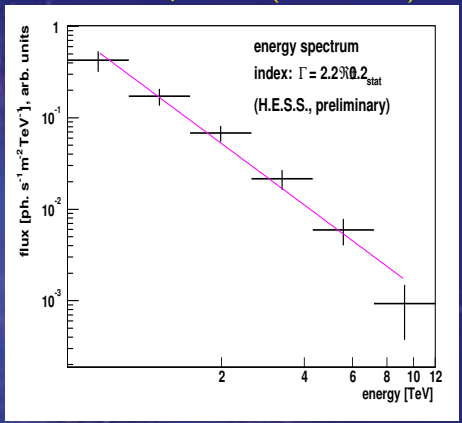


HESS J1303-63: The 2^{nd} unidentified TeV-source

Extended object ($\approx 0.2^\circ$)



Powerlaw spectrum ($\alpha = 2.2 \pm 0.2$)

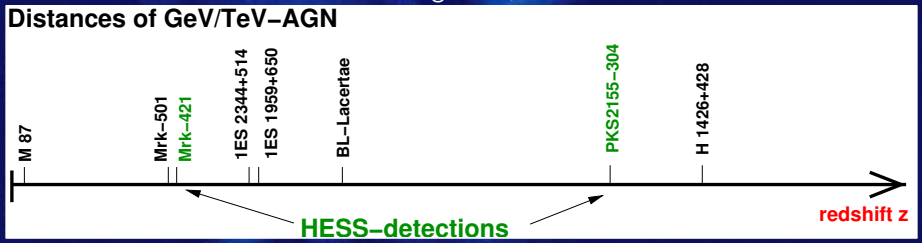


Populated corner of the sky but ...
No obvious radio / optical / X-ray counterpart found

$\longrightarrow$ Further HESS & Future GLAST observations

Extragalactic Sources

Physics: Production mechanisms, understanding the AGN family, Extragalactic Background, ...

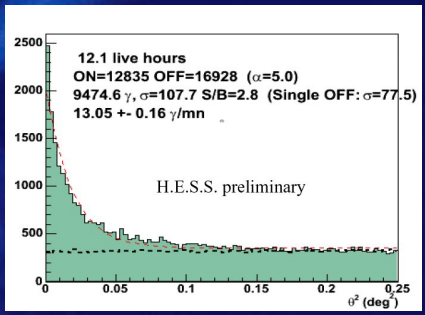


Mrk 421: HESS Detection

First Detection: Whipple (Punch et al. 1992), subsequently confirmed by many

HESS observations

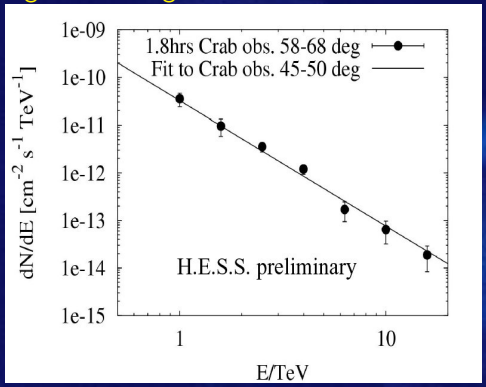
- 13.6 h observation full 4-telescope-system
- Average zenith angle: 63 deg
- Low state: 01/2004 - 6σ (2.1 h)
- High state: 04/2004 - $\approx 100\,\sigma$
  11.5 h, 11 γ/minute
- More than 8000 γ-ray events



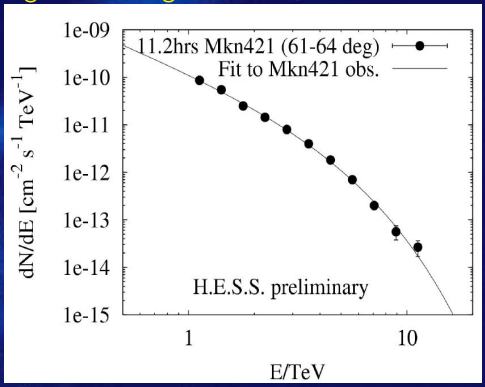
See contributions to GAMMA 2004 by D. Horns & A. Lemière

Mrk 421: Energy Spectrum

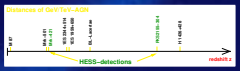
High zenith angles: Crab



High zenith angles: Mrk 421



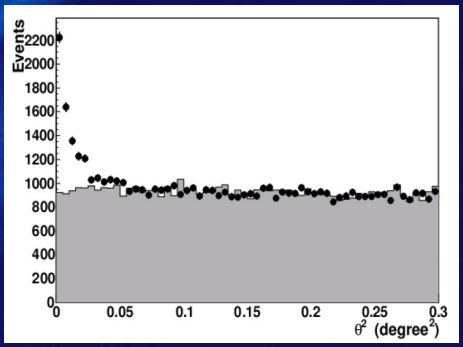
PKS 2155-304



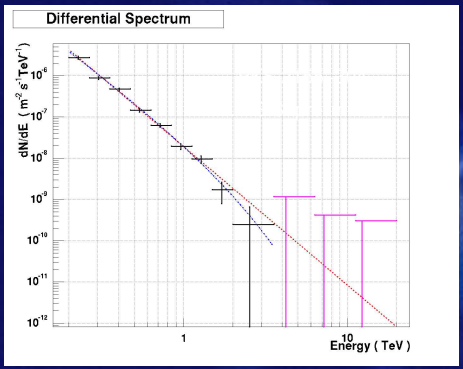
- High Frequency Peaked BL-Lac Object
- First detection: Durham Mark 6 Telescope
- $z = 0.116 \longrightarrow$ second most distant GeV/TeV emitter so far

HESS Data

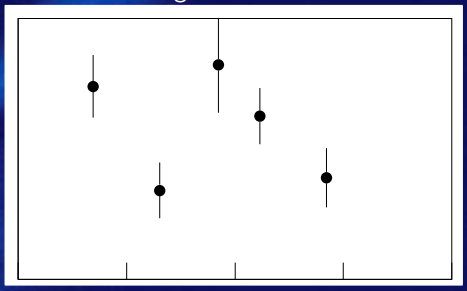
- Observation time $\approx$60 h
- Different data sets: varying Energy threshold 165 GeV - 305 GeV
- Significance: $>50\,\sigma$
- Many photons collected ($>$4000)
- Flux level consistent with Durham Mark6 detection
- No strong flux variations observed



PKS 2155-304: Energy Spectrum & Flux Variation



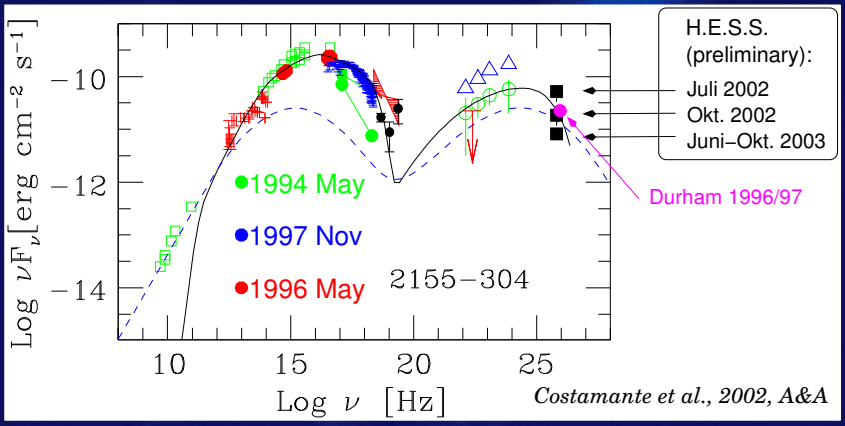
Intra-night flux variation:



No strong flux variations observed

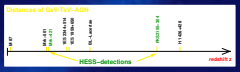
Object is detected *every night* $\longrightarrow$ Observation of a ≪ground state≫(first ever!)

PKS 2155-304: SSC model

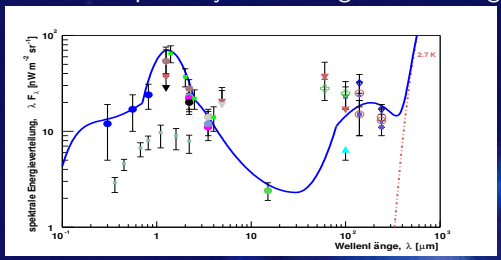


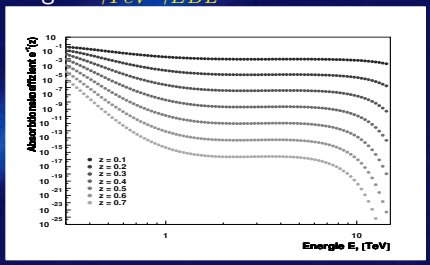
One-zone homogeneous SSC-model, Costamante et al. 2002

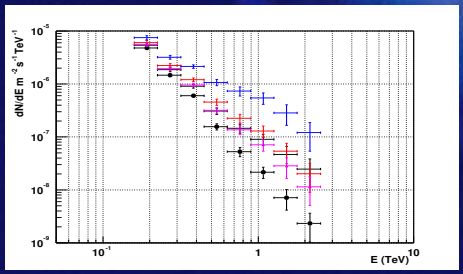
PKS 2155-304 & EBL Absorption



Absorption by the Extragalactic Background Light: $\gamma_{TeV}\,\gamma_{EBL} \longrightarrow e^+e^-$





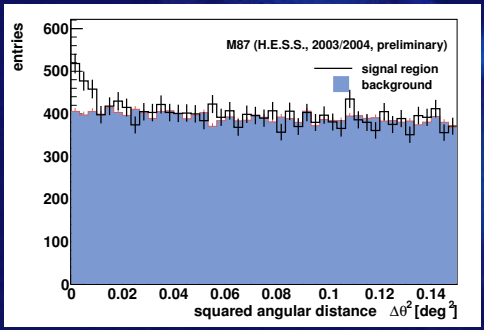


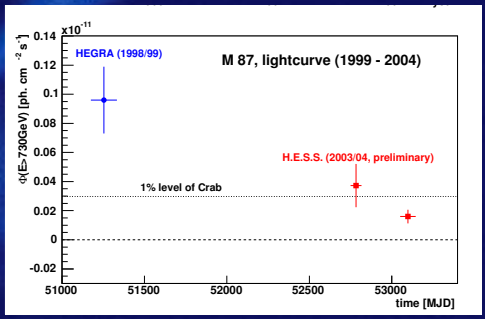
← PKS 2155-304 by HESS

- $z = 0.116$
- Several EBL models tried
- No cutoff seen yet
- Wait for higher statistics (2004) at high energies (?)

The Radio Galaxy M87

- distance 16 Mpc (z≈0.004)
- central BH : 2-3×10^9 $M_\odot$
- first evidence for TeV emission : HEGRA, 4 σ-level ($\approx$ 80 h)
- 45 h H.E.S.S.-data : $> 4\sigma$



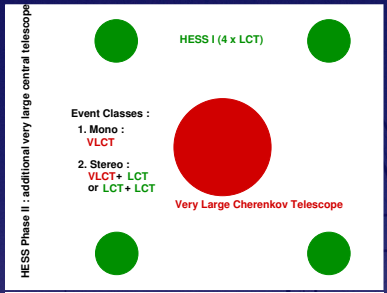


M. Beilicke et al., Texas Symposium

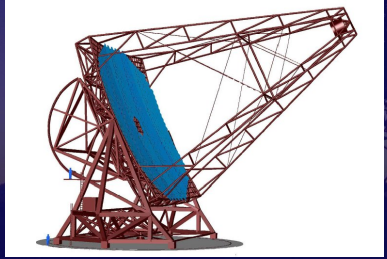
weakest GeV/TeV-source so far

The Near Future: H.E.S.S.Phase 2

HESS Phase 2 = HESS 1 + Very Large Cherenkov Telescope



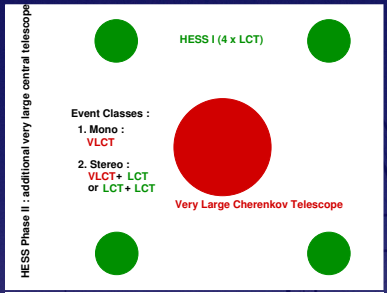
MAN Design: conventional alt-az mount



Very Large Cherenkov Telescope:

- Reflector : 28 m ($\approx 600\,m^2$)
- Focal distance ≈ 35 m
- Camera: diam. 2.5 m (≈ 2000 kg)
- 2048 PMTs (0.07°/pixel)
- FoV : 3.5°
- Trigger rate 2-20 kHz
- Faster ARS memories needed
- Minimize data flow: 2^{nd} level trigger

HESS Phase 2 = HESS 1 + Very Large Cherenkov Telescope



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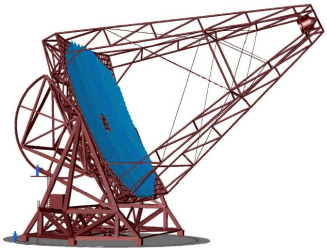
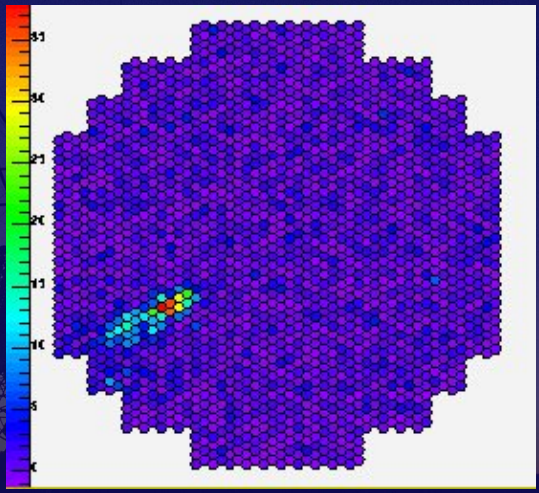


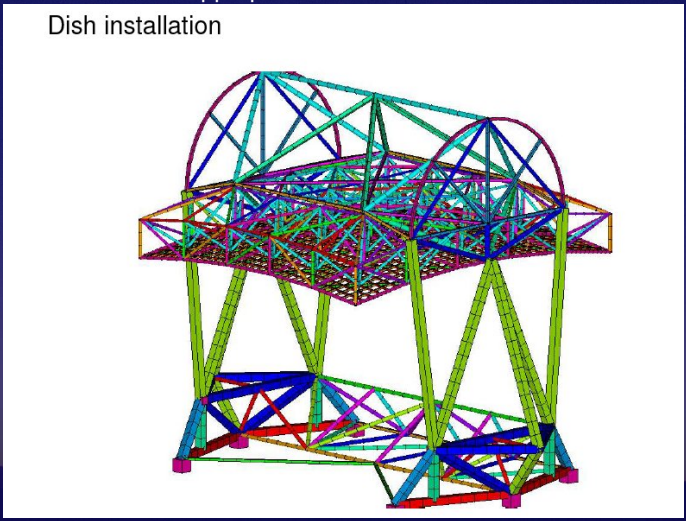
Simulation: 30 GeV γ



HESS Phase 2 in Namibia

Dish installation

No appropriate crane in Namibia ...



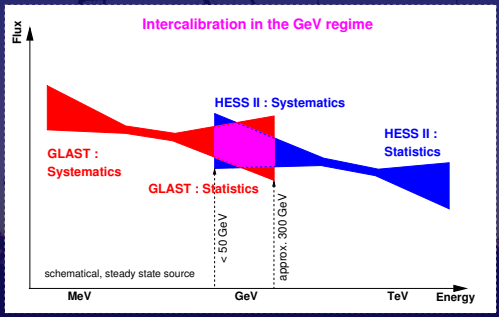
Motivation

- **Lower energy threshold / cover wider energy range**
  - $\longrightarrow$ Pulsars : cutoff at low energies / soft sectra
  - $\longrightarrow$ AGNs at large z : less absorbtion at low energies
  - $\longrightarrow$ other sources EGRET (...GLAST) : cutoff
- **complementarity with GLAST**
  - $\longrightarrow$ time resolution of H.E.S.S.
  - $\longrightarrow$ large field of view of GLAST

Objective : operation in best possible conditions
$\longrightarrow$ lowest threshold & highest γ count rates

Future collaboration with GLAST

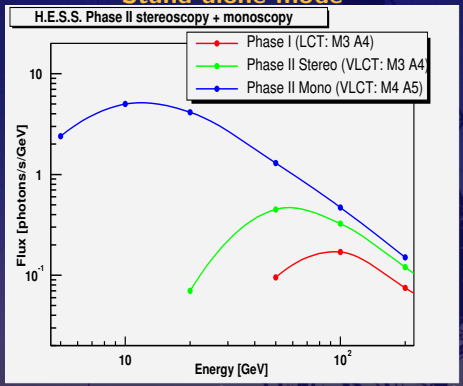
- Overlapping energy regime : Observations of the same particle population
  - GLAST trigger for HESS-observations
  - HESS will produce sensitive variability studies
- Simultaneous observations of a steady source $\longrightarrow$ Intercalibration



Performance of the stand alone VLCT without L2 Trigger

MC simulation : 4 LCTs + VLCT

Stand alone mode



- monoscopy
- L1 : $N_{pix} = 4$, $A_{pix} = 5$
- no L2 trigger !

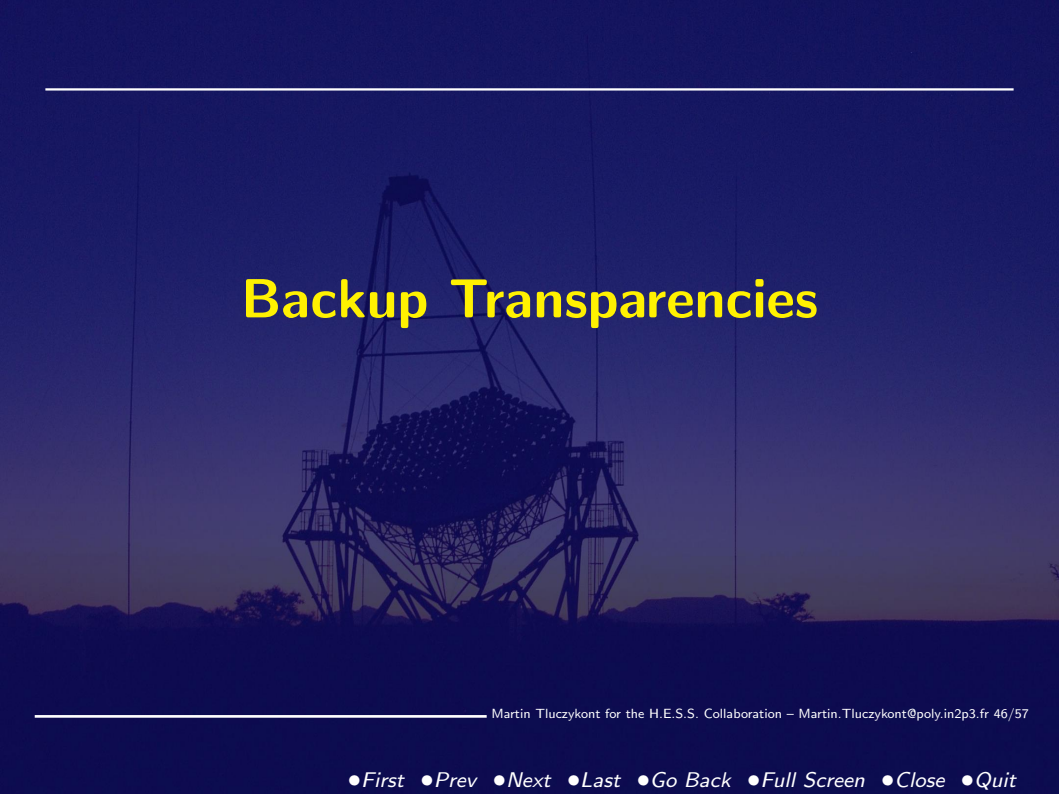
- $E_{seuil} \approx 20\,GeV$
- # events $\gamma \times 20$ (peak)
- trigger rate : 2.5 kHz
- dead time $<< 5\,\%$

$\Longrightarrow$ improvement using stereo + mono VLCT

further possible improvement $\Longrightarrow$ Lower PM thresholds

Summary

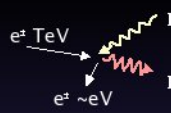
- HESS Phase I completed and fully operational Planned performance achieved
- Outstanding new results
  - 1^{st} astronomical GeV/TeV image (RX J1713.7)
  - New object class (Psr B1259-63)
  - 1^{st} time two sources in one FoV
  - 1^{st} observation of BL Lac at low state? (PKS 2155-304)
- $\longrightarrow$ soon more...
- HESS Phase II is being planned:
  - Objective: low energies & improvement of current system
  - Method: additional Very Large Central Telescope
  - Towards 10-30 GeV



Backup Transparencies

Gamma Ray Emission Models

TeV gamma ray production via Inverse Compton Scattering of the star's photon field and accelerated particle populations from:

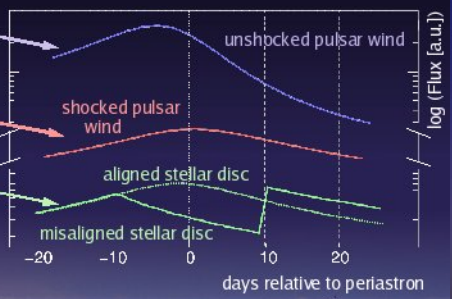


5) Unshocked pulsar wind *Ball, Kirk* 2000

6) Pulsar wind termination shock induced by star's photon field *Ball, Dodd* 2001

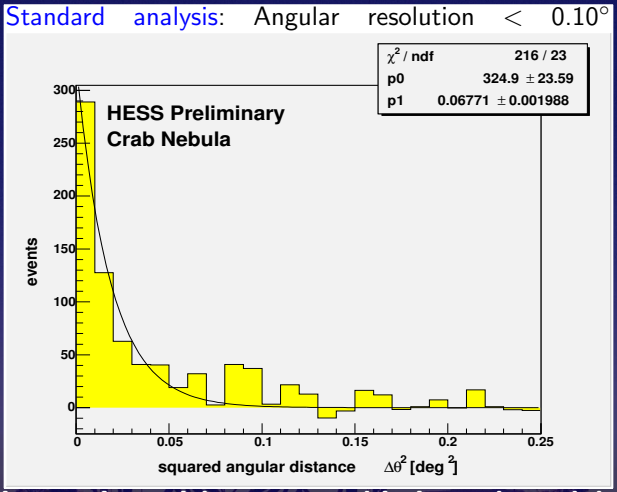
7) Interaction region of star mass outflow and pulsar wind *Kawachi et al.* 2004

All models predict flux variation dependent on orbital phase



Probe particle acceleration with star's photon field!

HESS 1 Performance: Angular resolution



Improving with more sophisticated models (Semi-Analytical model & 3D-Model)

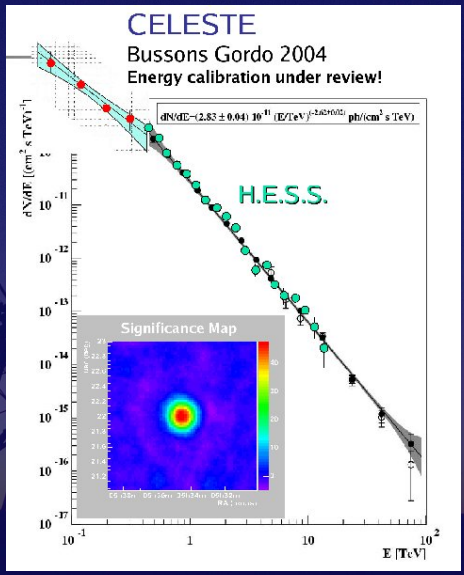
HESS 1 Performance: Energy Reconstruction

MC

- Energy resolution: 10-20 %
- Reconstruction bias $< 10\,\%$

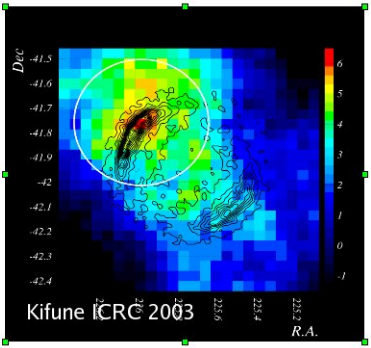
Crab Nebula Spectrum

- $\alpha = 2.59 \pm 0.04^{stat}$
  - Whipple: $\alpha = 2.49 \pm 0.07$
  - HEGRA: $\alpha = 2.62 \pm 0.05$
  - CAT: $\alpha = 2.76 \pm 0.04$
- Flux (E>1 TeV): $2.16 \pm 0.08 \times 10^{-7}\ \mathrm{m}^{-2}\mathrm{s}^{-1}$
  - Whipple: $2.15^{\pm 0.43} \cdot 10^{-7}\ \mathrm{m}^{-2}\mathrm{s}^{-1}$
  - HEGRA: $1.75^{\pm 0.37} \cdot 10^{-7}\ \mathrm{m}^{-2}\mathrm{s}^{-1}$
  - CAT: $1.78^{\pm 0.39} \cdot 10^{-7}\ \mathrm{m}^{-2}\mathrm{s}^{-1}$

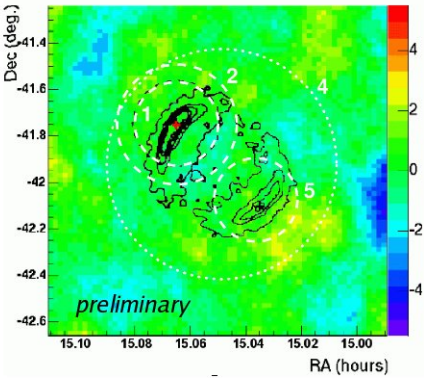


The shell-type Supernova Remnant SN 1006: Inconsistency

CANGAROO SN 1006



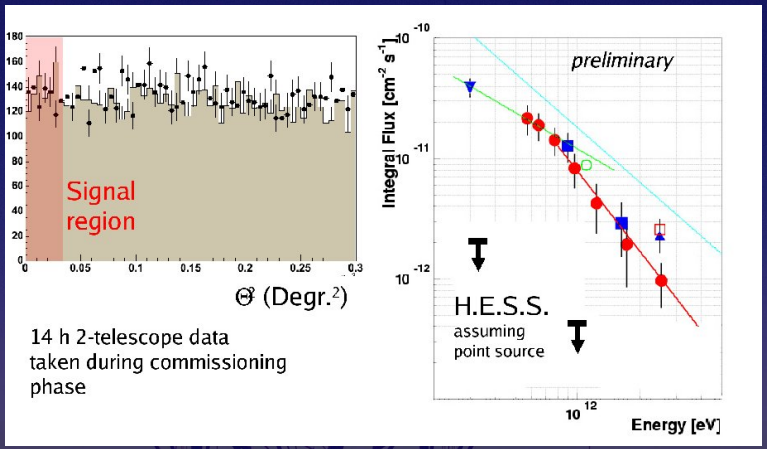
H.E.S.S. significance map



Tanimori et al., Ap J 497 (1988) L25
3.8 m telescope
+ conference proceedings

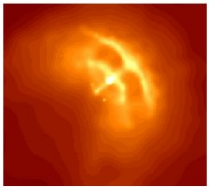
Hofmann, Gamma 2004

PSR B1706-44 by HESS



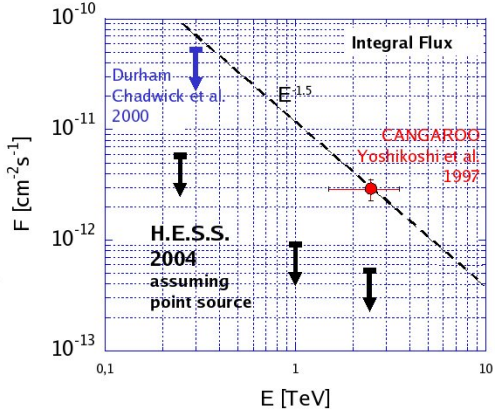
Hofmann, Gamma 2004

Vela by HESS



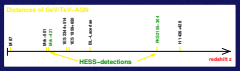
Chandra

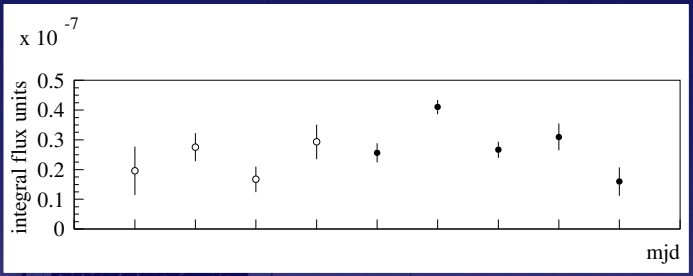
- CANGAROO source 0.13° off pulsar
- H.E.S.S. limits similar for both CANGAROO and pulsar location

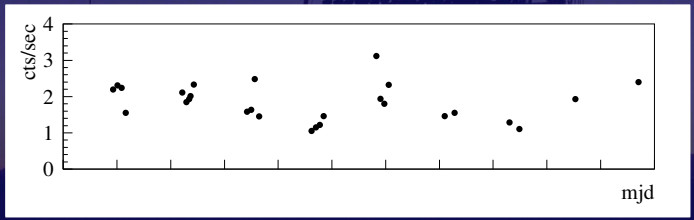


Hofmann, Gamma 2004

PKS 2155-304: Multi Wavelength Campaigns



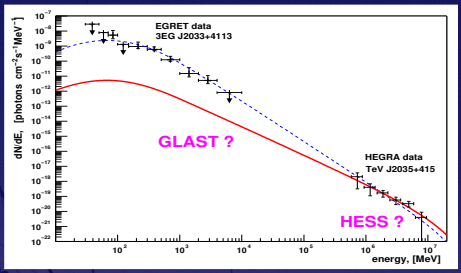


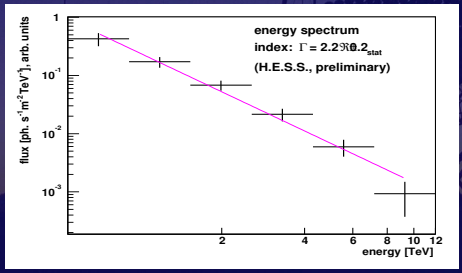


Multiwavelength campaigns carried out, results in preparation (B. Giebels)

What kind of objects to observe for inter calibration?

- Steady state source: simplifies a lot
- Hadronic
  - Naive: straight power law
  - More efficient acceleration at GeV energies ?
- Leptonic ?



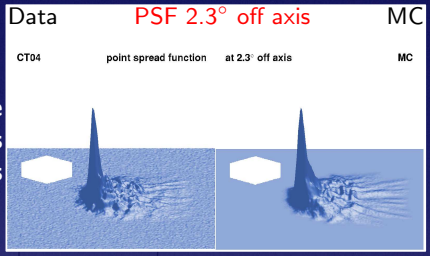


Possible Candidate: HESS 1303-63

- Power law: $\alpha = 2.2$
- Flux: 0.1 Crab (E>1 TeV)
- GLAST-detection: O(days)

The Mirror Alignment

Individual mirror facets steerable
Alignment of facets using stars
PSF well within specifications



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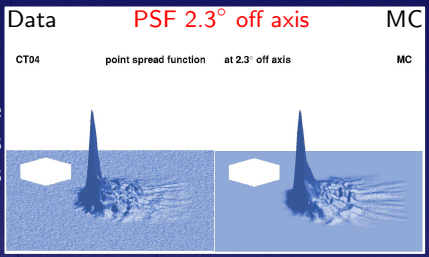

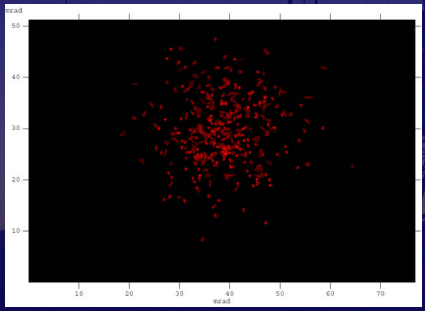

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